

Black-hole microstate spectroscopy: ringdown, quasinormal modes, and echoes

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Deep conceptual problems associated with classical black holes can be addressed in string theory by the “fuzzball” paradigm, which provides a microscopic description of a black hole in terms of a thermodynamically large number of regular, horizonless, geometries with much less symmetry than the corresponding black hole. Motivated by the tantalizing possibility to observe quantum gravity signatures near astrophysical compact objects in this scenario, we perform the first $3+1$ numerical simulations of a scalar field propagating on a large class of multicenter geometries with no spatial isometries arising from $\mathcal{N} = 2$ four-dimensional supergravity. We identify the prompt response to the perturbation and the ringdown modes associated with the photon sphere, which are similar to the black-hole case, and the appearance of echoes at later time, which is a smoking gun of the absence of a horizon and of the regular interior of these solutions. The response is in agreement with an analytical model based on geodesic motion in these complicated geometries. Our results provide the first numerical evidence for the dynamical linear stability of fuzzballs, and pave the way for an accurate discrimination between fuzzballs and black holes using gravitational-wave spectroscopy.

I. INTRODUCTION

Within Einstein’s theory of General Relativity, black holes (BHs) are the simplest macroscopic objects one can conceive. In stationary configurations, they are fully described only by their mass, spin, and possibly electric charge [1–4], being in this respect more akin to elementary particles than to astrophysical objects [5]. This simplicity is also associated with a high degree of symmetry: stationary BHs must be axisymmetric [6], and become spherical in the static (i.e., non-spinning) limit.

Owing to this and other remarkable properties, the equations governing the linear response of a BH to external perturbations and its quick relaxation towards stationarity after being formed (e.g., in a merger or in a stellar collapse) are separable in terms of a simple set of ordinary differential equations [7, 8], which enormously simplifies the analysis of BH linear perturbations. The latter are crucial, for instance, to describe the so-called “ringdown” during the post-merger phase of a binary coalescence [9]. The BH ringdown is governed by a discrete set of complex frequencies – the so-called quasinormal modes (QNMs) – which are uniquely determined by the BH parameters. BH spectroscopy [10–18] performed by measuring the ringdown with current and future gravitational-wave detectors [19–21] is at present the most robust way to study the strong-field regime of General Relativity and the nature of a merger remnant [19, 20, 22, 23]. Indeed, BHs can be considered as the “hydrogen atom” of gravity, and their gravitational-wave spectrum is a unique footprint of possible deviations from General Relativity in the strong-field regime, similarly to the energy levels of the hydrogen whose measurement had a paramount impact in shaping the devel-

opment of quantum electrodynamics [24].

This state of affairs is enormously more involved when the spacetime fails to be as simple and as symmetric as a BH. This happens arguably in any quantum gravity proposal aiming at resolving some outstanding issues with classical BHs [25], namely the curvature singularities that are conjectured to be always covered by event horizons [26–28], the conundrum of the huge BH entropy [29, 30], and the unitarity-loss problem associated with Hawking evaporation at the semiclassical level [31].

In the string-theory “fuzzball” proposal [32–35], a classical BH is described by an ensemble of smooth and horizonless geometries which represent the microstates of the BH with the same mass and asymptotic charges. The classical properties of a BH are expected to emerge either through an averaging procedure over a large number of microstates or as a ‘collective behavior’ of fuzzballs. Although finding a statistically significant number of microstate geometries in order to account for the whole BH entropy is challenging, large families of microstates have been discovered in the last few years [36–41]. A microscopic description of the whole entropy is provided by D-brane counting for BPS black holes in four and five dimensions [42–44]. Should the fuzzball program be successful, it would be a natural solution to the singularity, entropy, and unitarity problems that plague the classical BH interpretation.

Fuzzball microstates¹ are much less symmetric than a BH. Besides being stationary solutions to consistent

¹ With some abuse of language we interchangeably use the terms “fuzzballs” and “microstates” (or “microstate geometries”), although as we said above the former are ensembles of the latter.

low-energy truncations of string theory, fuzzballs do not generically possess any spatial isometry. This lack of symmetry and the complexity of the microstate geometries have so far hampered the possibility to study their ringdown and multipole structure [45–53] and to compare it with the one of a BH, which is a task of utmost importance to devise phenomenological tests of the fuzzball paradigm – and hence of quantum gravity – with gravitational-wave data.

In order to overcome this problem, in this work we perform for the first time 3 + 1 numerical simulations of small fields propagating on a large class of a microstate geometries. We unveil the entire ringdown phenomenology [54, 55] predicted in less motivated models of exotic compact objects. This includes the universal prompt ringdown similar to the BH case, which is nonetheless followed by a modulated series of repeated “echoes” [56, 57] associated with long-lived modes almost trapped within the fuzzball gravitational potential, providing a smoking gun of a horizonless compact object [55]. These ringdown features are currently searched for in gravitational-wave data [23, 58–65] (see [54, 66] for some reviews). Finally, we provide a simple physical interpretation of these effects in terms of the geodesics and multipolar structure of these complicated microstate geometries. We use natural units throughout.

II. FUZZBALL GEOMETRIES

Our framework is that of $\mathcal{N} = 2$ four-dimensional supergravity, wherein gravity is (non-)minimally² coupled to four $U(1)$ gauge fields and three complex scalars. Microstates of spherically symmetric 4-dimensional BHs can be constructed in this theory by considering a multi-center system of intersecting D3-branes [40, 67–69]. The metric reads

$$ds^2 = -e^{2U}(dt + \omega)^2 + e^{-2U} \sum_{i=1}^3 dx_i^2, \quad (1)$$

with

$$e^{-4U} = Z_1 Z_2 Z_3 V - \mu^2 V^2, \quad (2)$$

$$*_3 d\omega = \frac{1}{2} (V dW - W dV + K^I dL_I - L_I dK^I),$$

and

$$Z_I = L_I + \frac{|\epsilon_{IJK}| K^J K^K}{2V}, \quad (3)$$

$$\mu = \frac{W}{2} + \frac{L_I K^I}{2V} + |\epsilon_{IJK}| \frac{K^I K^J K^K}{6V^2} \quad (4)$$

where ϵ_{IJK} is the totally antisymmetric tensor, and $\{V, L_I, K^I, W\}$ are eight harmonic functions ($I = 1, 2, 3$).

² Gauge kinetic functions and Kähler metric of the scalars are non-canonical.

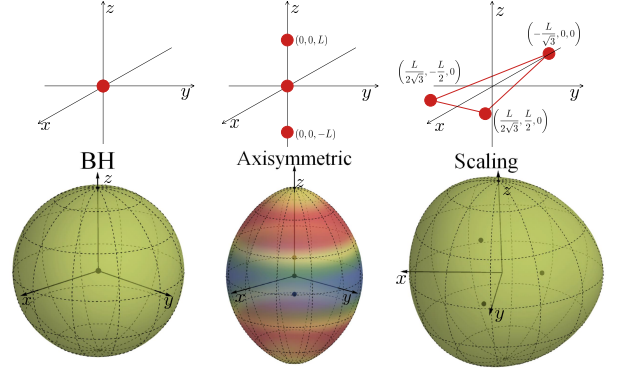


FIG. 1. Top: Schematic representation of the multicenter microstate solutions considered in this work (see Appendix A for details on their construction). The leftmost BH solution corresponds to the case in which the centers coincide. The middle solution is axially and equatorially symmetric, whereas the rightmost “scaling” solution breaks the axial symmetry. Bottom: corresponding embedding diagram for each solution. The shape describes constant- t and constant- r surfaces of the metric deep down the fuzzball’s throat. Deformations are related to the mass multipole moments, whereas the colors are weighted according to $g_{t\phi}$ to represent the leading current multipole moments (which vanish for the BH and the scaling solution) [52, 70].

As an ansatz, we take N -center harmonic functions of the form

$$V = 1 + \sum_{a=1}^N \frac{v_a}{|\vec{x} - \vec{x}_a|}, \quad L_I = 1 + \sum_{a=1}^N \frac{\ell_{I,a}}{|\vec{x} - \vec{x}_a|}, \quad (5)$$

$$K^I = \sum_{a=1}^N \frac{k_a^I}{|\vec{x} - \vec{x}_a|}, \quad W = \sum_{a=1}^N \frac{m_a}{|\vec{x} - \vec{x}_a|}, \quad (6)$$

with $\vec{x} = r(\sin \theta \cos \phi, \sin \theta \sin \phi, \cos \theta)$. From a four-dimensional perspective these microstate geometries are singular at the centers but, for specific choice of charges and positions of the centers, they admit a regular horizonless five-dimensional uplift. Overall, the solution is regular, horizonless, and free of other pathologies (e.g., closed timelike curves). In general it carries four electric $Q_A = (Q_0, Q_I)$ and four magnetic $P^A = (P^0, P^I)$ charges. For concreteness we shall focus on 3-center ($N = 3$) solutions and restrict ourselves to 4-charge solutions by imposing $Q_0 = P^I = 0$, although the generalization is straightforward. We provide their explicit form in Appendix A. These metrics have a finite non-vanishing entropy and asymptotically reduce to the 4-charge BPS BH solution [71] which generalizes the extremal Reissner-Nordström BH to the case of four different charges and three scalar fields. In the isotropic coordinates, the line element of the 4-charge BH reads [71]

$$ds^2 = -f(r)dt^2 + f(r)^{-1} [dr^2 + r^2(d\theta^2 + \sin^2 \theta d\phi^2)], \quad (7)$$

with $f(r) = (H_1 H_2 H_3 H_4)^{-1/2}$ and $H_A = 1 + \frac{Q_A}{r}$. The total mass is $M = \frac{1}{4} (Q_1 + Q_2 + Q_3 + Q_4)$ (with $Q_4 = P^0$ in the cases we consider). When $Q_A = Q$, the solution reduces to the extremal Reissner-Nordström BH with mass $M = Q$.

Although our method is general, for concreteness we shall consider some representative class of microstate geometries, depicted in Fig. 1. The first family comprises axisymmetric solutions with equatorial symmetry, i.e. reflection symmetric with respect to the equatorial plane, $\theta \rightarrow \pi - \theta$ (equivalently, $z \rightarrow -z$). They have three centers aligned along the z -axis, located at $\vec{x}_a = (0, 0, z_a)$, with

$$z_1 = L, \quad z_2 = 0, \quad z_3 = -L. \quad (8)$$

This is a discrete three-parameter family of solutions with nonvanishing angular momentum, whose explicit form is provided in Appendix A 2. The size L is determined in terms of the charges of the centers.

The second family of 3-center solutions we consider are those with $v_a = 1$, four charges, and centers located on the vertices of a triangle. This is a five-parameter family which in general breaks both axial and equatorial symmetry [52, 70]. The simplest element is the so-called “scaling” solution that corresponds to the three centers being located at the vertices of an equilateral triangle and has two free parameters – related to the mass and the size of the triangle – while the angular momentum vanishes. The explicit form of this solution is given in Appendix A 3 a.

In all cases, when the centers collapse to a single point, spherical symmetry is restored and the solution reduces to the extremal (non-rotating) BH.

III. LINEAR RESPONSE OF BHS AND MICROSTATES: PROMPT RINGDOWN, QNMS, AND ECHOES

We now focus on the linear dynamics of a neutral massless scalar field propagating in the metrics described in the previous section. This is governed by the Klein-Gordon equation $\square\Phi(t, r, \theta, \phi) = 0$. Decomposing Φ in spherical harmonics, i.e. $\Phi = \sum_{lm} \tilde{\Phi}_{lm}(r, t) Y_{lm}(\theta, \phi)$, the (orbital) angular-momentum number is an integer $l = 0, 1, 2, \dots$ and the azimuthal number m is an integer such that $|m| \leq l$. When the metric is spherically symmetric (as in the BH case considered above), the azimuthal number is $(2l+1)$ -fold degenerate and different l modes are decoupled from each other. In the axisymmetric case, the degeneracy of m is broken but m is still a conserved quantum number: modes with different m are decoupled. In the general case with no isometries, modes with different l and m can all mix with each others and it is more convenient to solve the Klein-Gordon equation directly as a $3+1$ evolution problem.

The prompt response of a compact object to some perturbation is universally described by the resonant exci-

tation of its photon sphere, where (unstable) closed null orbits reside. For a BH spacetime, the photon-sphere modes coincide with the QNMs of the object and dominate the linear response [56]. The QNMs are defined in the frequency domain as those complex eigenfrequencies, $\omega = \omega_R + i\omega_I$, which correspond to a solution that satisfies purely outgoing-wave boundary conditions at infinity and “regularity” conditions at the inner boundary. In the BH case, regularity at the horizon imposes purely ingoing-wave boundary conditions, whereas in the horizonless fuzzball case Φ must be regular at the origin $r = 0$. For each (l, m) mode, there is a countably infinite number of QNMs identified by the overtone index $n = 0, 1, 2, \dots$, with $n = 0$ labelling the fundamental QNM dominating the linear response at late times.

Crucially, if the spacetime is horizonless, the photon-sphere modes dominate only the initial ringdown until the perturbation has time to probe the inner boundary. Following the universal prompt ringdown, radiation can be reflected back and get quasi-trapped within the gravitational potential of the object, occasionally tunnelling to infinity and producing a series of repeated and modulated “echoes” [56, 57, 72, 73]. This transient regime interpolates between the prompt ringdown and the very late-time behavior, which is instead dominated by the long-lived modes of the horizonless compact object [55].

A. Analytical results in the geodesic approximation

There exists a tight relation between the ringdown modes of a compact object and some geodesic properties associated with the existence of an unstable photon sphere, as established in the eikonal limit [74, 75]. In the static case, the real part of the QNM frequency is related to the (azimuthal) orbital frequency, whereas the imaginary part of the QNM corresponds to the Lyapunov exponent of the orbit [75]. In the rotating case the relation between modes with generic (l, m) and specific geodesic quantities is more involved [76]. Strictly speaking the geodesic approximation is valid when $l \gg 1$ but it often works remarkably well also for smaller values of l [77].

Geodesic motion for a massless neutral particle moving in the spacetime given in Eq. (1) can be described by the null Hamiltonian $\mathcal{H} = \frac{1}{2} g^{\mu\nu} P_\mu P_\nu = 0$, where P_μ is the particle four-momentum. For simplicity we focus on the axisymmetric case, i.e. consider a stationary metric as in Eq. (1), where $\omega = \omega_\phi d\phi$. The Hamiltonian in this case can be written as

$$2\mathcal{H} = -e^{-2U} E^2 + e^{2U} \left(P_r^2 + \frac{P_\theta^2}{r^2} + \frac{(P_\phi + \omega_\phi E)^2}{r^2 \sin^2 \theta} \right) \quad (9)$$

where P_ϕ and $E = -P_t$ are constants of motion, while P_r, P_θ vary along the trajectory. We notice that even assuming axial symmetry, ω_ϕ and U typically depend both on r and θ , so in general the radial and angular

κ	M	L/M	J/M^2	r_+/M	r_-/M	$M\omega_{\text{QNM},+}^{\text{fuzzball}}$	$M\omega_{\text{QNM},-}^{\text{fuzzball}}$	$M\omega_{\text{QNM}}^{\text{BH}}$	$\Delta t_+/M$	$\Delta t_-/M$
3	13.75	0.1091	0.0714	0.7314	0.6759	$0.6827 - 0.0767i$	$0.7324 - 0.0878i$	$0.6996 - 0.0871i$	40.7	40.1
4	24.25	0.0508	0.0544	0.7596	0.6490	$0.6860 - 0.0841i$	$0.7372 - 0.0855i$	$0.7083 - 0.0859i$	63.8	62.7
5	37.75	0.0301	0.0439	0.7493	0.6379	$0.6919 - 0.0852i$	$0.7380 - 0.0839i$	$0.7128 - 0.0851i$	84.9	83.7
6	54.25	0.0201	0.0367	0.7378	0.6345	$0.6966 - 0.0854i$	$0.7373 - 0.0831i$	$0.7154 - 0.0847i$	105.1	104.1
7	73.75	0.0144	0.0315	0.7283	0.6343	$0.7001 - 0.0853i$	$0.7362 - 0.0827i$	$0.7170 - 0.0844i$	125.0	124.0
8	96.25	0.0109	0.0276	0.7207	0.6352	$0.7028 - 0.0852i$	$0.7351 - 0.0825i$	$0.7180 - 0.0841i$	144.6	143.7
9	121.75	0.0085	0.0246	0.7146	0.6367	$0.7050 - 0.0851i$	$0.7341 - 0.0825i$	$0.7188 - 0.0840i$	164.0	163.2
10	150.25	0.0069	0.0221	0.7097	0.6383	$0.7067 - 0.0849i$	$0.7332 - 0.0824i$	$0.7193 - 0.0839i$	183.3	182.6
50	3750	0.0003	0.0044	0.6746	0.6593	$0.7188 - 0.0837i$	$0.7244 - 0.0830i$	$0.7216 - 0.0834i$	942.2	942.1
100	15000	0.0001	0.0022	0.6706	0.6629	$0.7103 - 0.0835i$	$0.7231 - 0.0832i$	$0.7217 - 0.0833i$	1887.3	1887.2

TABLE I. Summary of the ringdown features of an axisymmetric fuzzball geometry (with $\kappa_1 = \kappa_2 = \kappa_3 = \kappa$) in the geodesic approximation for $l = m = \pm 2$. M is the total mass, the ratio L/M characterizes the distance of the centers, J/M^2 is the dimensionless angular momentum. The \pm signs refers to co-rotating and counter-rotating orbits. r_{\pm}/M are the dimensionless critical radii. For the BH case, the QNMs are in agreement with an exact frequency-domain computation (see Appendix B 2) within a few percent.

dynamics cannot be disentangled in simple terms. The situation improves if one further assumes equatorial symmetry. For this choice $\dot{P}_\theta = -\partial_\theta \mathcal{H} = 0$ at $\theta = \pi/2$, and a particle initially moving along the equator will remain on the plane.

The null Hamiltonian condition $\mathcal{H} = 0$ follows from (9) after setting $P_\theta = 0$ and $\theta = \pi/2$. One finds

$$P_r^2 - Q(r) = 0 \quad (10)$$

with radial effective potential

$$Q(r) = -\frac{1}{r^2} [P_\phi - b_+(r)E] [P_\phi - b_-(r)E] \quad (11)$$

and impact parameter functions

$$b_{\pm}(r) = -\omega_\phi(r) \pm r e^{-2U(r)}. \quad (12)$$

A particle falling from infinity will evolve according to Eq. (10) till it reaches a turning point r_* , i.e. a zero of $Q(r)$, and then bounces back to infinity. If the inversion point r_* is a double root of $Q(r)$, the point cannot be reached in a finite time and the particle gets trapped forever orbiting around the mass center and approaching asymptotically a circular orbit (the light ring). This happens for a critical choice E_c of the energy and of the radius r_c obtained by solving the critical equations

$$Q(E_c, r_c) = \partial_r Q(E_c, r_c) = 0. \quad (13)$$

Using Eq. (11), these equations can be written in the simple form

$$b'_c(r_c) = 0 \quad , \quad E_c = \frac{P_\phi}{b_c(r_c)} \quad (14)$$

where a prime denotes the radial derivative. Rotation produces an inner and an outer photon sphere radius

r_{\pm} obtained as the extrema of the $b_{\pm}(r)$ functions. The two signs distinguish between co-rotating and counter-rotating orbits with respect to the angular momentum of the spinning microstate geometry.

The QNM frequencies can be extracted from the WKB formula [50, 53, 74, 75]

$$\frac{Q(r_c)}{\sqrt{2Q''(r_c)}} = i \left(n + \frac{1}{2} \right), \quad (15)$$

after replacing E by the complex number $E = \omega_R + i\omega_I$ with a small imaginary part ω_I . This equation can be solved at leading order in ω_I by taking $\omega_R = E_c$, with E_c given by (14). Expanding then to linear order in ω_I and plugging into (15) one finds

$$\omega_{\text{QNM}} = \frac{l + \frac{1}{2}}{|b_c|} - i\lambda_c(2n + 1), \quad (16)$$

with $l = |m| = |P_\phi|$, for equatorial geodesics, and

$$\lambda_c = \left(\frac{\partial Q(r_c)}{\partial E} \right)^{-1} \sqrt{\frac{Q''(r_c)}{2}}. \quad (17)$$

The above result simplifies considerably in the case of spherical symmetry. For example, for the 4-charge BH with $Q_1 = Q_2 = M$, $Q_3 = Q_4 = \beta^2 M$, we get³

$$\omega_{\text{QNM}}^{\text{BH}} M = \frac{l + \frac{1}{2}}{(1 + \beta)^2} - i(2n + 1) \frac{\sqrt{\beta/2}}{(1 + \beta)^3}. \quad (18)$$

Interestingly, in the case of regular, horizonless objects the geodesic approximation can also capture the relevant timescales associated with echoes [56–58, 72, 73, 78].

³ For $\beta = 0$, one has a 2-charge system with zero horizon area ('small BH').

Geodesics with impact parameter $b = l/E$ along the equatorial plane are described by the radial equation

$$v_r(r, b) = \frac{dr}{dt} = \frac{\frac{\partial \mathcal{H}}{\partial P_r}}{\frac{\partial \mathcal{H}}{\partial P_t}} \approx \frac{2\sqrt{Q(r)}}{\frac{\partial Q(r)}{\partial E}}, \quad (19)$$

so the scattering time is given by

$$\Delta t(b) = 2 \int_{r_c}^{r_t} \frac{dr}{v_r(r, b)} \quad (20)$$

with r_t the turning point of the orbit. The time delay as a function of b has a minimum near $b = 0$. Minimal time delay provides an estimate of the delay between two subsequent echoes in the fuzzball geometry.

In Table I we provide a summary of the ringdown features in the geodesics approximation for a representative microstate geometry and compare them with the corresponding BH case. Although the method is generically valid for any axisymmetric spacetime, for concreteness we focus on the 3-parameter family of axisymmetric fuzzball solutions presented in the previous section with $\kappa_1 = \kappa_2 = \kappa_3 = \kappa$, which is regular for any integer $\kappa \geq 2$. Note that (equatorial) light rings exist in this geometry only when $\kappa \geq 3$.

Some comments are in order. First, we note that the distance between the centers of the microstate geometry, L/M , monotonically decreases as κ grows. Correspondingly the solution approaches the BH limit and the fuzzball photon-sphere QNMs coincide with those of the corresponding BH with the same mass. It is also interesting to note that the BH case ($L \rightarrow 0$ or $\kappa \rightarrow \infty$) maximizes the real part of the QNM. This feature is analogous to the fact that the Lyapunov exponent of unstable null geodesics near the photon sphere is maximum for certain BH solutions [50]. Furthermore, as $\kappa \gg 1$ the time delays Δt_{\pm} grows, showing that gravitational time dilation inside the fuzzball becomes larger in this limit. Based on these results, we would expect that the prompt ringdown of a fuzzball should be very similar to that of a BH, but extra features in the ringdown should appear on a timescale Δt_{\pm} , which is typically much longer than the decay time, $\sim -1/\omega_I$, of the fundamental BH QNM. In the next section we shall confirm these expectations by comparing the analytical approximation with fully numerical simulations.

B. Numerical setup for time evolution

Here, we explain the numerical set up for the time evolution of a scalar field in a microstate geometry. Owing to its spherical symmetry, the case of a 4-charge BH is much simpler and can be studied with a $1+1$ evolution code or in the frequency domain, as discussed in Appendix B. However, since the general fuzzball spacetime does not have spatial isometries, the field equation for the scalar field is not separable. We therefore evolve the scalar field

using $3+1$ dimensional numerical simulations. The numerical implementation is based on the Einstein Toolkit infrastructure [80–82] with mesh refinement provided by the Carpet package [83, 84], and multipatch infrastructure provided by Llama [85]. The scalar field is evolved in the ScalarEvolve code, which has been previously used and tested in [86–88]. We implement a 4-th order Runge-Kutta method, and the spatial derivatives are evaluated by 4-th order finite differences. The buffer zones between different refinement levels are evaluated by using 5-th order interpolation in space and 2-nd order interpolation in time.

Compared to the standard case of fields evolving on a BH spacetime, more resolution layers are needed. Indeed, the fuzzball spacetime has high curvature regions around the centers which must be resolved during the evolution and, at the same time, the asymptotic behavior of the field must be accurately extracted at large distances in order to get the ringdown signal. Furthermore, although regular from the five dimensional perspective, four-dimensional fuzzball geometries are singular at the centers and such singular behavior should be regularized to perform stable numerical simulations. We regularize the geometry by replacing

$$\frac{1}{|\vec{x} - \vec{x}_a|} \rightarrow \frac{\text{erf}(|\vec{x} - \vec{x}_a|/\epsilon)}{|\vec{x} - \vec{x}_a|} \quad (21)$$

in the harmonic functions V , L_I , K^I , and W . The error function smoothly interpolates between unity when $|\vec{x} - \vec{x}_a| \gg \epsilon$ and $\propto |\vec{x} - \vec{x}_a|$ when $|\vec{x} - \vec{x}_a| \ll \epsilon$. Furthermore, the expressions of the one-form ω has a coordinate singularity on the axis for axisymmetric fuzzball solutions (see Eq. (A10)), which can be regularized by the following replacement⁴:

$$\frac{1}{1 - (\vec{n}_a \cdot \vec{n}_{ab})^2} \rightarrow \frac{1}{1 - (\vec{n}_a \cdot \vec{n}_{ab})^2 + \epsilon}. \quad (22)$$

Details about code testing and the regularization procedure of spacetime are given in Appendix C.

Concerning the initial data, we have performed a variety of simulations with different initial profiles. We report here two representative cases. The first type is a spherically symmetric Gaussian ‘shell’:

$$\Phi(0, x) = A e^{-\left(\frac{r-r_0}{\sigma}\right)^2}, \quad (23)$$

whereas the second type is an $l = m = 2$ profile:

$$\Phi(0, x) = \mathcal{W}(r, r_{\max}, r_{\min}) \text{Re}(Y_{2,2}(\theta, \phi)) \frac{r}{\sigma} e^{-\frac{r-r_0}{\sigma}}, \quad (24)$$

where σ and r_0 define the typical width and central value location of the initial profile, and $\mathcal{W}(r)$ is a window function, which is a 5-th order smooth polynomial satisfying

⁴ Using two patches with $\omega_{\pm} = \omega \pm d\phi$ is not practical in numerical analyses.

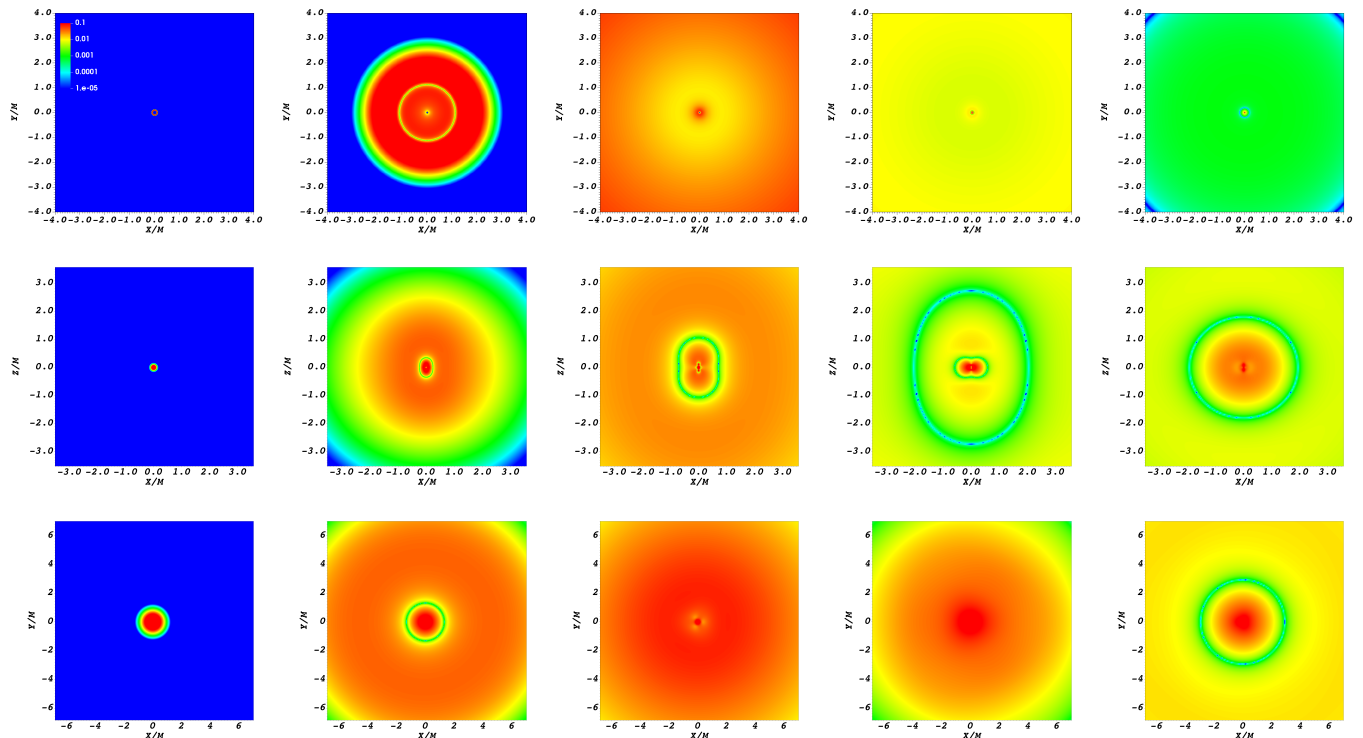


FIG. 2. Snapshots of the evolution of a scalar field around the geometries depicted in Fig. 1 and studied in the following, namely: 4-charge BH (top panel) corresponding to left panel in Fig. 3, an axisymmetric fuzzball (middle panel) corresponding to Fig. 5, and a 3-center scaling microstate (bottom panel) corresponding to bottom panel in Fig. 7. From left to right, they are snapshots at $t/M = 0, 15, 30, 45, 60$. Related movies are publicly available [79].

$\mathcal{W}(r_{\min}, r_{\max}, r_{\min}) = 0$ and $\mathcal{W}(r_{\max}, r_{\max}, r_{\min}) = 1$. In both cases the initial scalar field is instantaneously static, i.e.

$$(\partial_t - \gamma^i \partial_i) \Phi(0, x) = 0, \quad (25)$$

where γ^i is the shift vector in the $3+1$ decomposition of the metric [89].

IV. NUMERICAL RESULTS

Here, we summarize the numerical results and discuss the scattering of a scalar wavepacket off fuzzball geometries. Figure 2 shows some representative snapshots of our simulations on the $y = 0$ plane. Related movies are publicly available [79].

A. Ringdown of 4-charge BHs

We evolve the scalar field around a 4-charge BH with $Q_1 = Q_3$ and $Q_2 = Q_4$ based on the formulation discussed in Appendix B, and extract the field at fixed radius. Figure 3 shows the evolution of the scalar field with $l = 0$ and $l = 1$. We observe the standard QNM ringing, wherein the signal dies off exponentially until producing

the typical late-time power-law tail due to backscattering off the gravitational potential [90]. Figure 4 shows the QNM frequency as a function of the single dimensionless parameter of the metric, $\beta = \sqrt{Q_3/Q_1}$, with $\beta \rightarrow 1$ being the extremal Reissner-Nordström limit. The data points in Fig. 4 are extracted from the simulations by fitting the time-domain waveform, whereas the dotted and solid curve respectively show the QNMs given by the analytical geodesic approximation [Eq.(18)] and by an exact frequency-domain code (see Appendix B 2). The analytical geodesic result is in excellent agreement with the exact one (especially for large l), and both are in good agreement with the results extracted from the time domain. The agreement of the imaginary part of the $l = 0$ mode is less good (but still within a few percent) due to the smaller number of ringdown cycles before the power law (compare left and right panels of Fig. 4), which reduces the accuracy of the fit. Since the BH QNMs extracted from the time evolution are in good agreement with the photon sphere QNMs computed with the geodesic approximation, and since the latter are independent of the object interior, we expect the prompt ringdown to be described by the photon sphere modes also in the fuzzball case [56, 57]. In next subsection, we shall confirm this expectation and use it to distinguish the photon-sphere QNMs from other characteristic modes of the system.

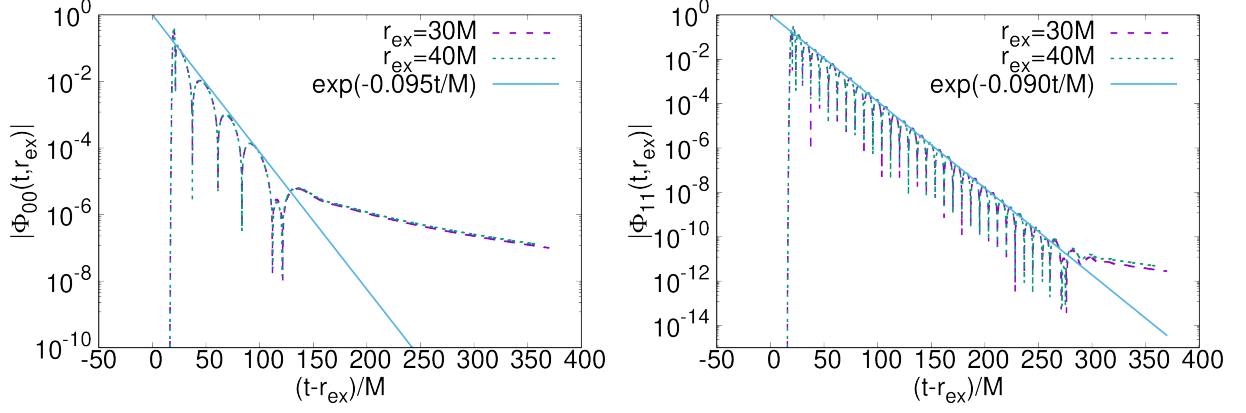


FIG. 3. Time evolution of $(l, m) = (0, 0)$ (left) and $(l, m) = (1, 1)$ (right) multipole mode of a scalar field around a 4-charge BH solution with $(Q_1 = Q_3, Q_2 = Q_4) = (1.1, 0.9)M$ at two different extraction radii. For the initial data, we took a instantaneously static Gaussian centered at $1.1M$ in tortoise coordinates and with width $\sigma = 0.01M$ (see Appendix B). The general behavior is independent of the initial data.

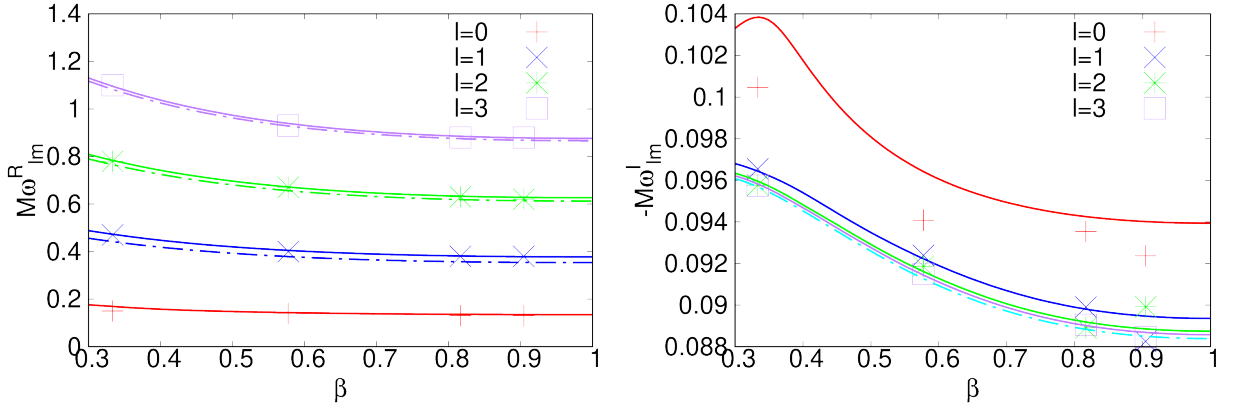


FIG. 4. Real part (left panel) and imaginary part (right panel) of the fundamental QNMs of a 4-charge BH with $Q_1 = Q_3$ and $Q_2 = Q_4 = \beta^2 Q_1$ as a function of β for each l . Solid and dot-dashed curves refer to the exact frequency-domain computation (see Appendix B 2) and the analytical result of the geodesic analysis (see Eq. 18), respectively. Data points are the modes extracted from the time-domain waveform of our simulations.

B. Ringdown and echoes of fuzzball microstates

Let us start by discussing the case of axisymmetric fuzzball solutions. Figure 5 shows the time evolution for initial data given in Eq. (23) with $\sigma = 0.036M$ and $r_0 = 0$ around a $\kappa_1 = \kappa_2 = \kappa_3 = \kappa = 3$ axisymmetric fuzzball, taken as a representative example (the qualitative properties do not change for $\kappa > 3$). The absence of spherical symmetry produces mode mixing: the initially spherical profile excites both spherical ($l = m = 0$) and quadrupolar ($l = 2, m = 0$) modes, where all dipolar ($l = 1$) modes and the quadrupolar modes with $m \neq 0$ vanish identically due to the symmetries of the metric and to the angular momentum composition rules. Hexadecupolar ($l = 4, m = 0$) modes are also excited but are subleading and not shown in the plot.

In the prompt-ringdown phase of both modes, we observe the standard damped oscillations with a decay rate

of about $0.1M$, in agreement with the geodesic result for the photon-sphere modes (see Table I and Fig. 4). At very late times, the decay is much slower and presumably corresponds to the long-lived QNMs of the fuzzball geometry. At intermediate times, we observe the typical echoes expected for horizonless geometries [56–58, 72, 73]. Echoes are more evident for the “induced” quadrupolar modes compared to the spherical mode. The reason for this is twofold: i) the effective potential for modes with large angular momentum is higher, so generically modes with larger l are more efficiently confined; ii) the fuzzball’s gravitational potential has a quadrupolar component that can trap $l = 2$ modes more efficiently, especially for mostly equatorial $l = m = 2$ modes discussed below. Therefore, for spherical modes, echoes could be absent or possibly buried in the low-frequency fuzzball fundamental QNM that dominates the late response. For $l = 2, m = 0$, the typical time delay between echoes

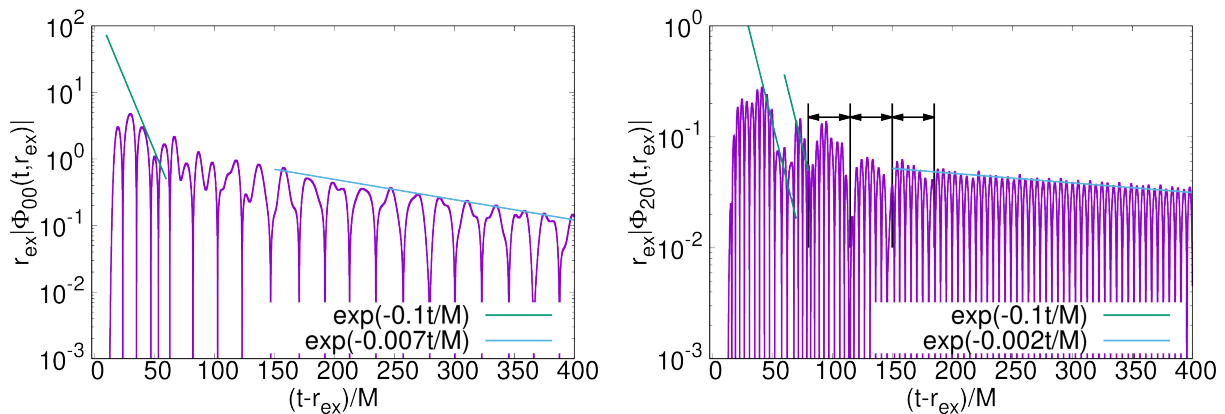


FIG. 5. Time evolution of $l = m = 0$ (left panel) and $l = 2, m = 0$ (right panel) multipole modes of the scalar field around axisymmetric fuzzball solution with $\kappa_1 = \kappa_2 = \kappa_3 = \kappa = 3$, and extracted at $r_{\text{ex}} = 23M$. The evolution starts from instantaneously static spherically symmetric Gaussian profile [Eq. (23)] with $\sigma = 0.036M$ and $r_0 = 0$.

is roughly $\Delta t = 35M$. Note, however, that it is difficult to identify accurately the echo delay time from the time evolution, especially due to the presence of multiple pulses. Indeed, our instantaneously static Gaussian pulse initially splits into an outgoing pulse and an incoming pulse; the latter gets quickly reflected at $r = 0$ and follows the outgoing pulse with a short delay [91], which is approximately $\Delta T \approx 16M$, as computed in the geodesic approximation for this solution. Therefore, we expect a complicate echo pattern from the modulation of this doublet.

Note also that the echo delay time in Fig. 5 cannot be directly compared with Δt_{\pm} in Table I, since the latter is valid for equatorial motion that corresponds to $l = m$ modes.⁵ In order to perform such a comparison, in Fig. 6 we show the evolution starting from $l = 2, m = 2$ initial data [Eq. (24)]. At early times, we observe the QNMs of the photon sphere (with a short decay time denoted by the green line), which is in good agreement with the estimation from the geodesic approximation and is indeed similar to the BH QNMs. After the prompt ringdown, the expected echo pattern appears. The delay time is highlighted by black arrows and is approximately $40M$, also in good agreement with the time scale predicted by the geodesic analysis. Some further interesting examples are presented in Appendix D.

Finally, let us focus on generic 3-center microstate solutions. Although our method is general, for concreteness we focus on the scaling solution wherein the centers form an equilateral triangle (Fig. 1). This solution is equatorially symmetric but breaks axial symmetry so the geodesic approximation discussed in Sec. III A does not apply and we have to rely on numerical results only.

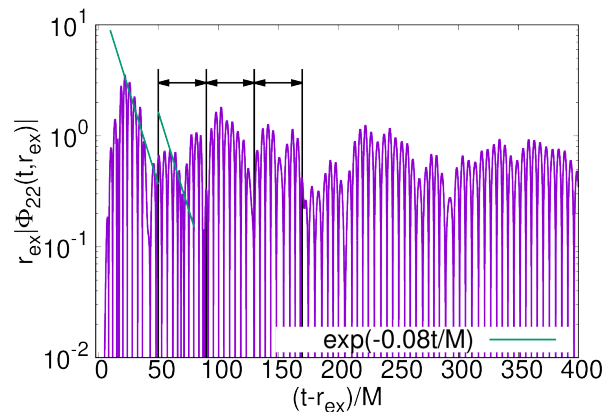


FIG. 6. Same as in Fig. 5 but for the evolution of quadrupolar $l = m = 2$ perturbations with initial data described by an instantaneously static $l = 2, m = 2$ Gaussian profile [Eq. (24)] with $\sigma = 0.036M$, $r_0 = 0.14M$, $r_{\text{max}} = 0.07M$, and $r_{\text{min}} = 0.04M$. The extraction radius is $r_{\text{ex}} = 12M$. The prompt ringdown and the echo delay time are in agreement with the predictions in Table I.

Figure 7 shows that the evolution of the scalar field starting from the spherically symmetric, instantaneously static, Gaussian initial data around two representative examples of scaling solution. The ringdown of this family of solutions shows the same qualitative features as in the axisymmetric case, in particular the echoes are more evident in the “induced” quadrupolar⁶ mode $l = 2, m = 0$,

⁵ However, preliminary analysis suggests that the timescale for $m \neq l$ geodesic motion is similar to the $l = m$ case. Therefore, the echo time scale of left panel in Fig. 5 is compatible with the analytical approximation.

⁶ Note that, even if the solution is not axisymmetric, the $l = 2, m = \pm 1$ modes are not excited by an initially spherical pulse. The reason is that the scaling solution only has $l = 2, m = 0$ and $l = 3, m = \pm 3$ multipole moments or higher [52, 70]. According to the standard angular-momentum sum rules, these multipoles cannot produce a mode with $m = 1$ when coupled to a spherical ($l = m = 0$) initial pulse.

whereas the spherical mode displays immediately the long-lived, low-frequency fuzzball QNMs, that have a much longer decay time relative to the corresponding photon-sphere quantity. Although not shown, the evolution of $l = m = 2$ initial data is similar to that presented in Fig. 6 for the axisymmetric solution.

V. DISCUSSION

We have developed a general technique to study the linear response of a large family of regular, horizonless fuzzball geometries describing the microstates of extremal spherically-symmetric BHs. These microstates do not possess spatial isometries, which makes the problem dramatically more involved than in the BH case.

For the first time, we studied the evolution of scalar perturbations on these geometries by using a $3 + 1$ numerical-relativity code. Our method is generic and can be applied to any stationary fuzzball. We unveiled the whole ringdown phenomenology studied in recent years for exotic compact objects [54, 55]; in particular, we showed that the prompt ringdown of fuzzballs is associated with the photon-sphere modes and mimics the BH response, whereas the presence of echoes [56–58, 72, 73] at late time is a smoking gun of the absence of a horizon and of the presence of the fuzzball’s regular throat.

In all our simulations, the perturbations always decay in time, providing strong numerical evidence that the fuzzball microstates under consideration are linearly stable, at least until our maximum simulation time, $t \approx 1000M$. In addition, these geometries do not possess an ergoregion, so they are not plagued by the ergoregion instability of horizonless compact objects [92–98].

Although in viable astrophysical scenarios BHs are expected to be neutral, charged BH solutions to supergravity are a useful toy model to explore the properties of their corresponding microstates. The numerical analysis developed here does not rely on supersymmetry and can be directly applied to putative non-supersymmetric and neutral microstates, should the latter be found. In this case, we expect that all the features found in this work will be qualitatively the same.

A natural extension of our work is to consider other types of perturbations. While adapting our code to study a test vector field or any self-interacting (possibly massive) test fields is straightforward, considering metric perturbations is technically possible but much more challenging, since in these charged geometries the perturbations of the metric, gauge fields, and scalar fields are coupled to each other. However, also in this case we expect that the salient features of the ringdown will remain qualitatively the same.

A more promising avenue could be to embed this family of microstate geometries into an effective model such as the one based on the membrane paradigm developed in [55]. Our results can be used to calibrate the model and check if it reproduces the salient features of the

fuzzball ringdown phenomenology. An important missing ingredient in current models is the fact that – owing to the lack of spatial symmetries – different angular modes can be excited with comparable amplitude, even for highly symmetric initial data, as shown by our simulations.

On the longer run, the goal is to confront the predictions of the fuzzball scenario with gravitational-wave data, for example extending current ringdown tests [20, 23, 55] and echo searches [23, 58–65] to include the predictions of this model, which are dramatically more complex than the toy models used so far to model gravitational-wave echoes.

On the theory side, our results also urge to address the “measure problem” with fuzzballs. As we have shown, in general the ringdown of a single microstate geometry is very peculiar and contains clear features that can be used to distinguish it from the standard BH ringdown. However, if a classical BH is described by a quantum superposition of microstates, then macroscopic observables (e.g. QNM frequencies and damping times) might be related to an average over the entire ensemble. Whether the signals should sum incoherently (possibly destroying the echo pattern) or what is the phase space and measure of this average remain open problems that should be urgently addressed if one wishes to test quantum-gravity effects near compact objects in the gravitational-wave era.

Studies of possible observational signatures of the fuzzball scenario are just in their infancy, but are acquiring considerable attention in the context of gravitational-wave and electromagnetic tests [49, 51, 52, 70] that may also encode string corrections [99–101] to memory effects [102]. We hope that our study on BH microstate spectroscopy could pave the way to further developments in this exciting field.

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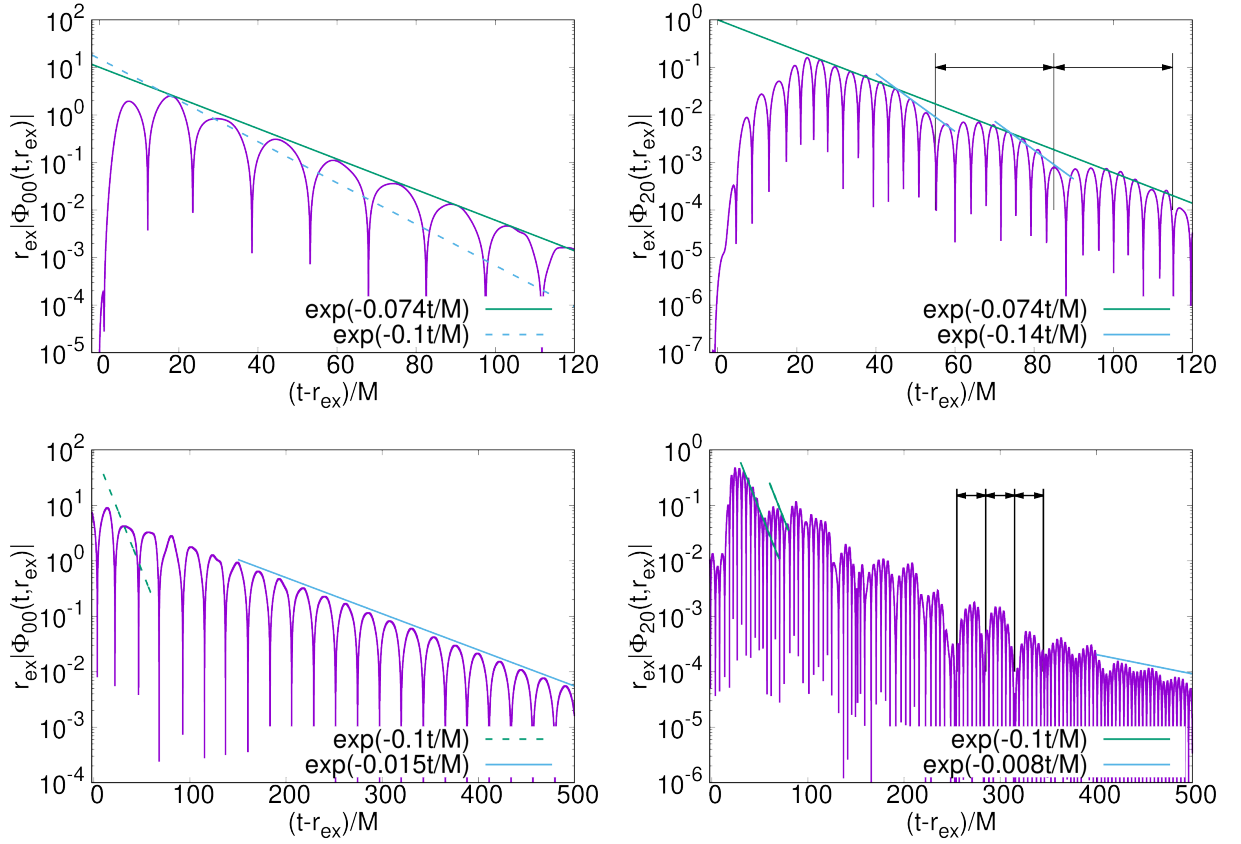


FIG. 7. Time evolution of $l = 0, m = 0$ (left panel) and $l = 2, m = 0$ (right panel) multipole modes of the scalar field around scaling solution with $L = 0.67M, \kappa = 1$ (top panel), and $L = 0.27M, \kappa = 2$ (bottom panel). The evolution starts from instantaneously static spherically symmetric Gaussian profile [Eq.23] with $\sigma = 0.67M$ (top panel) and $\sigma = 0.27M$ (bottom panel). The extraction radius is $r_{\text{ex}} = 20.0M$ (top panel) and $r_{\text{ex}} = 8.0M$ (bottom panel)

mental physics.” M.B., D.C., A.G. and J.F.M. would like to acknowledge partial support by Grant ID1202 “Strong

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Appendix A: Fuzzball microstate geometries

In this appendix we provide details on the fuzzball microstate geometries considered in the main text.

1. General multicenter fuzzball solutions

We are interested in solutions of the form (1) admitting a regular horizonless five-dimensional uplift with metric

$$ds_5^2 = -(Z_1 Z_2 Z_3)^{-\frac{2}{3}} [dt + \mu(d\Psi + w_0) + w]^2 + (Z_1 Z_2 Z_3)^{\frac{1}{3}} [V^{-1}(d\Psi + w_0)^2 + V d\vec{x}^2] \quad (\text{A1})$$

and $*_3 dw_0 = dV$. It can be seen that the resulting metric reduces to \mathbb{R} times a Gibbons-Hawking space near the centers if μ vanishes and Z_I remain finite near the centers. The origin of the Gibbons-Hawking space can display orbifold singularities if $|v_a| \neq 1$ but the latter can be easily regularized by splitting the centers (see [67] for a review). The finiteness of Z_I is ensured by setting

$$\ell_{I,a} = -\frac{1}{2} \frac{|\epsilon_{IJK}| k_a^J k_a^K}{v_a}, \quad , \quad m_a = \frac{k_a^1 k_a^2 k_a^3}{v_a^2} \quad (\text{A2})$$

while the vanishing of μ leads to the so called “bubble equations” [67]

$$\sum_{b=1}^N \frac{\Gamma_{ab}}{r_{ab}} = \Lambda_a, \quad (\text{A3})$$

where $r_{ab} = |\vec{x}_a - \vec{x}_b|$,

$$\Gamma_{ab} = v_a v_b \prod_{I=1}^3 \left(\frac{k_a^I}{v_a} - \frac{k_b^I}{v_b} \right), \quad \Lambda_a = \sum_{I=1}^3 k_a^I - \frac{k_a^1 k_a^2 k_a^3}{v_a^2}. \quad (\text{A4})$$

Notice that Eq. (A3) gives $N - 1$ equations for the distances r_{ab} and one equation for the coefficients k_a^I , namely

$$\sum_{a=1}^N \Lambda_a = 0. \quad (\text{A5})$$

In order to remove singularities and closed time-like curves one should impose the following conditions

$$e^{2U} > 0, \quad Z_I V > 0. \quad (\text{A6})$$

Note that the first condition prevents the existence of an ergoregion and therefore these geometries are free of the ergoregion instability that exists for horizonless ultracompact objects [92–98] as well as for non-supersymmetric fuzzballs of the so-called JMaRT family [103–106].

The explicit expression for the one-form ω can be derived from Eq. (2) by making use of the bubble equations (A3) and relations (A2). One finds

$$\omega = \frac{1}{4} \sum_{a,b} \Gamma_{ab} \omega_{ab}, \quad (\text{A7})$$

with $\omega_{ab} = \vec{\omega}_{ab} \cdot d\vec{x}$ a solution of

$$\vec{\nabla} \times \vec{\omega}_{ab} = \frac{1}{|\vec{x} - \vec{x}_a|} \vec{\nabla} \frac{1}{|\vec{x} - \vec{x}_b|} - \frac{1}{|\vec{x} - \vec{x}_b|} \vec{\nabla} \frac{1}{|\vec{x} - \vec{x}_a|}. \quad (\text{A8})$$

The solution can be written as

$$\omega_{ab} = \frac{(\vec{n}_a + \vec{n}_{ab}) \cdot (\vec{n}_b - \vec{n}_{ab})}{r_{ab}} d\phi_{ab}, \quad (\text{A9})$$

with

$$\vec{n}_a = \frac{\vec{x} - \vec{x}_a}{|\vec{x} - \vec{x}_a|}, \quad \vec{n}_{ab} = \frac{\vec{x}_a - \vec{x}_b}{r_{ab}}, \quad d\phi_{ab} = \frac{(\vec{n}_a \times \vec{n}_{ab}) \cdot d\vec{x}}{|\vec{x} - \vec{x}_a| [1 - (\vec{n}_a \cdot \vec{n}_{ab})^2]}. \quad (\text{A10})$$

The solution carry in general four electric, $Q_A = (Q_0, Q_I)$, and four magnetic, $P_A = (P^0, P^I)$, charges respectively given by

$$\begin{aligned} P^0 &= \sum_a v_a, \quad Q_I = \sum_a \ell_{I,a} = -\frac{1}{2} \sum_a \frac{|\epsilon_{IJK}| k_a^J k_a^K}{v_a}, \\ P^I &= \sum_a k_a^I, \quad Q_0 = \sum_a m_a = \sum_a \frac{k_a^1 k_a^2 k_a^3}{v_a^2}. \end{aligned} \quad (\text{A11})$$

2. Axisymmetric three-center solutions with $v_a \neq 1$

We consider first a family of axisymmetric solutions with equatorial symmetry $z \rightarrow -z$ ($\theta \rightarrow \pi - \theta$). We align the three centers along the z -axis

$$\vec{x}_a = (0, 0, z_a) \quad , \quad z_1 = L \quad , \quad z_2 = 0 \quad , \quad z_3 = -L, \quad (\text{A12})$$

and take the charges to be

$$v_1 = v_3 = -1 \quad , \quad v_2 = 3 \quad , \quad k^I{}_a = \begin{pmatrix} -\kappa_1 & 0 & \kappa_1 \\ -\kappa_2 & 0 & \kappa_2 \\ -\kappa_3 & 0 & \kappa_3 \end{pmatrix} \quad (\text{A13})$$

This is therefore a 3-parameter family of solutions spanned by the integers κ_1 , κ_2 , and κ_3 . For this case, Γ_{ab} and Λ_a take the simple form

$$\Gamma_{ab} = \kappa_1 \kappa_2 \kappa_3 \begin{pmatrix} 0 & -\frac{v_2}{v_1^2} & -\frac{8}{v_1} \\ \frac{v_2}{v_1^2} & 0 & -\frac{v_2}{v_1^2} \\ \frac{8}{v_1} & \frac{v_2}{v_1^2} & 0 \end{pmatrix} \quad , \quad \Lambda_a = \kappa_1 \kappa_2 \kappa_3 \begin{pmatrix} -q + \frac{1}{v_1^2} \\ 0 \\ q - \frac{1}{v_1^2} \end{pmatrix} \quad (\text{A14})$$

with $q = \frac{\kappa_1 + \kappa_2 + \kappa_3}{\kappa_1 \kappa_2 \kappa_3}$. The bubble equations are solved for

$$L = \frac{4v_1 + v_2}{q v_1^2 - 1} \quad (\text{A15})$$

whereas the one-form $\omega = \omega_\phi d\phi$ takes the form

$$\omega_\phi = \sum_{b>a}^N \frac{\Gamma_{ab}}{2r_a r_b (z_a - z_b)} [(r_a + z_a - r \cos \theta)(r_b - z_b + r \cos \theta) - r^2 \sin^2 \theta] \quad (\text{A16})$$

with $N = 3$ and

$$r_a = |\vec{x} - \vec{x}_a| \quad (\text{A17})$$

so that

$$r_1^2 = r^2 + L^2 - 2rL \cos \theta \quad , \quad r_2^2 = r^2 \quad , \quad r_3^2 = r^2 + L^2 + 2rL \cos \theta \quad (\text{A18})$$

For completeness, we list the harmonic functions and other relevant quantities. The eight harmonic functions are

$$\begin{aligned} V &= 1 + v_1 \left[\frac{1}{r_1} + \frac{1}{r_3} \right] + \frac{v_2}{r_2} \quad , \quad W = \frac{\kappa_1 \kappa_2 \kappa_3}{v_1^2} \left[-\frac{1}{r_1} + \frac{1}{r_3} \right] \\ L_1 &= 1 - \frac{\kappa_2 \kappa_3}{v_1} \left[\frac{1}{r_1} + \frac{1}{r_3} \right] \quad , \quad L_2 = 1 - \frac{\kappa_1 \kappa_3}{v_1} \left[\frac{1}{r_1} + \frac{1}{r_3} \right] \quad , \quad L_3 = 1 - \frac{\kappa_1 \kappa_2}{v_1} \left[\frac{1}{r_1} + \frac{1}{r_3} \right] \\ K_1 &= -\kappa_1 \left[\frac{1}{r_1} - \frac{1}{r_3} \right] \quad , \quad K_2 = \kappa_2 \left[-\frac{1}{r_1} + \frac{1}{r_3} \right] \quad , \quad K_3 = \kappa_3 \left[-\frac{1}{r_1} + \frac{1}{r_3} \right] \end{aligned} \quad (\text{A19})$$

The nonvanishing charges are (remind $v_1 = v_3 = -1, v_2 = 3$)

$$P^0 = 2v_1 + v_2 \quad , \quad Q_1 = -\frac{2\kappa_2 \kappa_3}{v_1} \quad , \quad Q_2 = -\frac{2\kappa_1 \kappa_3}{v_1} \quad , \quad Q_3 = -\frac{2\kappa_1 \kappa_2}{v_1} \quad (\text{A20})$$

To ensure that the charges are positive, we assume

$$\kappa_1, \kappa_2, \kappa_3 \geq 1 \quad , \quad v_1 < 0 \quad , \quad v_2 + 2v_1 > 0 \quad (\text{A21})$$

The mass M and entropy S of the corresponding BH read

$$\begin{aligned} M &= \frac{P^0 + Q_1 + Q_2 + Q_3}{4} = \frac{v_1(v_2 + 2v_1) - 2(\kappa_1 \kappa_2 + \kappa_2 \kappa_3 + \kappa_1 \kappa_3)}{4v_1} \quad , \\ S^2 &\propto \lim_{r \rightarrow 0} r^4 e^{-4U} = -(v_2 + 2v_1) \frac{8\kappa_1^2 \kappa_2^2 \kappa_3^2}{v_1^3} \quad , \end{aligned} \quad (\text{A22})$$

and are positive by construction. The total angular momentum of the solution reads $\vec{J} = J \vec{e}_z$, with

$$J = -\kappa_1 \kappa_2 \kappa_3 \frac{4v_1 + v_2}{2v_1^2} \quad (\text{A23})$$

Notice that under the following scaling

$$v_a \rightarrow \gamma v_a \quad , \quad \kappa_i \rightarrow \gamma \kappa_i \quad , \quad (A24)$$

all the charges scale as

$$Q_I \rightarrow \gamma Q_I \quad , \quad P^0 \rightarrow \gamma P^0 \quad , \quad M \rightarrow \gamma M \quad , \quad J \rightarrow \gamma^2 J \quad , \quad S \rightarrow \gamma^2 S. \quad (A25)$$

This is an overall scaling, therefore dimensionless quantities such as $M\omega_{\text{QNM}}$ are not affected by it. This implies that any given solution can be scaled to a solution with arbitrarily large charges and mass.

Finally, the 4-charge BH associated to the fuzzball solution is obtained by bringing the three centers to the origin. In this case harmonic functions become

$$V = 1 + \frac{v_2 + 2v_1}{r} \quad , \quad L_1 = 1 + \frac{2\kappa_2\kappa_3}{v_1 r} \quad , \quad L_2 = 1 + \frac{2\kappa_1\kappa_3}{v_1 r} \quad , \quad L_3 = 1 + \frac{2\kappa_1\kappa_2}{v_1 r} \quad (A26)$$

$$K^I = W = 0 \quad ,$$

leading to

$$e^{-4U} = L_1 L_2 L_3 V \quad , \quad \omega = 0 \quad (A27)$$

The mass, charges, and entropy are given by Eqs. (A20)-(A22), whereas the BH is nonspinning, $J = 0$.

3. Three-center solutions with $v_a = 1$

The second family of solutions we consider are those with three centers, $v_a = 1$ and four charges $Q_0 = P^I = 0$. The general solution to the bubble equation is described by four integers κ_a specifying the k_a^I -matrix to be of the form

$$k_a^I = \begin{pmatrix} -\kappa_1\kappa_2 & -\kappa_1\kappa_3 & \kappa_1(\kappa_2 + \kappa_3) \\ \kappa_3 & \kappa_2 & -\kappa_2 - \kappa_3 \\ -\kappa_4 & \kappa_4 & 0 \end{pmatrix}. \quad (A28)$$

and therefore

$$\begin{aligned} V &= 1 + \sum_{a=1}^3 \frac{1}{r_a} \quad , \quad W = \kappa_1 \kappa_2 \kappa_3 \kappa_4 \left(\frac{1}{r_1} - \frac{1}{r_2} \right) \\ L_1 &= 1 + \kappa_4 \left(\frac{\kappa_3}{r_1} - \frac{\kappa_2}{r_2} \right) \quad , \quad L_2 = 1 + \kappa_1 \kappa_4 \left(-\frac{\kappa_2}{r_1} + \frac{\kappa_3}{r_2} \right) \\ L_3 &= 1 + \kappa_1 \left(\frac{\kappa_2\kappa_3}{r_1} + \frac{\kappa_2\kappa_3}{r_2} + \frac{(\kappa_2 + \kappa_3)^2}{r_3} \right) \quad , \quad K^1 = \kappa_1 \left(-\frac{\kappa_2}{r_1} - \frac{\kappa_3}{r_2} + \frac{\kappa_2 + \kappa_3}{r_3} \right) \\ K^2 &= \frac{\kappa_3}{r_1} + \frac{\kappa_2}{r_2} - \frac{\kappa_2 + \kappa_3}{r_3} \quad , \quad K^3 = \kappa_4 \left(-\frac{1}{r_1} + \frac{1}{r_2} \right) \end{aligned} \quad (A29)$$

again with $r_a = |\vec{x} - \vec{x}_a|$ and κ_i some arbitrary integers. The solution describes a microstate of a Reissner-Nordström BH with a magnetic charge P^0 and three electric charges Q_I given by

$$P^0 = 3 \quad , \quad Q_1 = \kappa_4(\kappa_3 - \kappa_2) \quad , \quad Q_2 = \kappa_1\kappa_4(\kappa_3 - \kappa_2) \quad , \quad Q_3 = \kappa_1(\kappa_2^2 + 4\kappa_2\kappa_3 + \kappa_3^2). \quad (A30)$$

The bubble equations constrain the distances $r_{ab} = |\vec{x}_a - \vec{x}_b|$ between the centers to be related by

$$\begin{aligned} r_{12} &= \frac{2\kappa_1\kappa_4(\kappa_2 - \kappa_3)^2 r_{23}}{\kappa_1\kappa_4(2\kappa_2^2 + 5\kappa_2\kappa_3 + 2\kappa_3^2) + (\kappa_2 + \kappa_4 - \kappa_1\kappa_3 + \kappa_1\kappa_2\kappa_3\kappa_4)r_{23}} \\ r_{13} &= \frac{\kappa_1\kappa_4(2\kappa_2 + \kappa_3)(\kappa_2 + 2\kappa_3)r_{23}}{\kappa_1\kappa_4(2\kappa_2^2 + 5\kappa_2\kappa_3 + 2\kappa_3^2) - (\kappa_1 - 1)(\kappa_2 + \kappa_3)r_{23}}. \end{aligned} \quad (A31)$$

a. Scaling solution

The simplest solution in this class is the so-called scaling solution, corresponding to the choice

$$\kappa_1 = 1 \quad , \quad \kappa_2 = 0 \quad , \quad \kappa_3 = \kappa_4 = \kappa . \quad (\text{A32})$$

For this choice one finds

$$P^0 = 3 \quad , \quad Q_I = (\kappa^2, \kappa^2, \kappa^2) \quad , \quad M = \frac{3(1+\kappa^2)}{4} \quad , \quad J = 0 . \quad (\text{A33})$$

In particular, note that the solution is non-spinning. The centers are located at the vertices of an equilateral triangle.

$$(x_1, y_1, z_1) = \left(\frac{L}{2\sqrt{3}}, \frac{L}{2}, 0 \right) , \quad (x_2, y_2, z_2) = \left(\frac{L}{2\sqrt{3}}, -\frac{L}{2}, 0 \right) , \quad (x_3, y_3, z_3) = \left(-\frac{L}{\sqrt{3}}, 0, 0 \right) . \quad (\text{A34})$$

b. Another axisymmetric solution

Although not considered in the main text, for completeness we provide here another axisymmetric solution. In this case axial symmetry is found for the choice $\kappa_2 = 0$ and $\kappa_4 = \kappa_1 \kappa_3$, i.e.

$$k^I_a = \kappa_3 \begin{pmatrix} 0 & -\kappa_1 & \kappa_1 \\ 1 & 0 & -1 \\ -\kappa_1 & \kappa_1 & 0 \end{pmatrix} . \quad (\text{A35})$$

The centers are located at

$$\vec{x}_a = (0, 0, z_a) \quad , \quad z_1 = L \quad , \quad z_2 = 0 \quad , \quad z_3 = -L , \quad (\text{A36})$$

with

$$L = \frac{\kappa_1^2 \kappa_3^2}{\kappa_1 - 1} . \quad (\text{A37})$$

The mass, charges, angular momentum and entropy of the solution read

$$\begin{aligned} P^0 &= 3, \quad Q_1 = \kappa_1 \kappa_3^2, \quad Q_2 = \kappa_1^2 \kappa_3^2, \quad Q_3 = \kappa_1 \kappa_3^2, \\ M &= \frac{3 + \kappa_1 \kappa_3^2(2 + \kappa_3)}{4}, \quad \vec{J} = -\frac{1}{2} \kappa_1^2 \kappa_3^3 \vec{e}_z \quad , \quad S = 3\kappa_1^4 \kappa_3^6 . \end{aligned} \quad (\text{A38})$$

Appendix B: Numerical implementation for 4-charge BH

1. Time-domain analysis

To evolve the scalar field in a spherically symmetric BH, it is convenient to define an effective potential for different angular modes. The isotropic coordinates in Eq. (7) can be transformed into areal radius coordinates as:

$$ds^2 = -f(r)dt^2 + g(r) \left(\frac{2\varrho}{r^2 g'(r) + 2r g(r)} \right)^2 d\varrho^2 + \varrho^2 d^2\Omega , \quad (\text{B1})$$

where $\varrho^2 = g(r)r^2$. In the general case of four independent charges, the analytical expression for $\varrho(r)$ is cumbersome. If we assume $Q_1 = Q_3$ and $Q_2 = Q_4$ for simplicity, then the horizon areal radius is $\varrho_H = \sqrt{Q_1 Q_2}$, and the metric reduces to

$$ds^2 = -B(\varrho)^2 dt^2 + A(\varrho)^2 d\varrho^2 + \varrho^2 d^2\Omega , \quad (\text{B2})$$

where

$$\begin{aligned} B(\varrho) &= \frac{\sqrt{(Q_1 - Q_2)^2 + 4\varrho^2} - (Q_1 + Q_2)}{2\varrho}, \\ A(\varrho) &= \frac{4\varrho^2}{(Q_1 - Q_2)^2 + 4\varrho^2 - (Q_1 + Q_2)\sqrt{(Q_1 - Q_2)^2 + 4\varrho^2}}. \end{aligned} \quad (\text{B3})$$

Since the space-time is spherically symmetric, we can separate the angular dependence of the field as

$$\Phi = \sum_{l,m} \frac{\sigma_{lm}(t, \varrho)}{\varrho} Y_{lm}(\theta, \phi), \quad (\text{B4})$$

where $Y_{lm}(\theta, \phi)$ are the spherical harmonics. The tortoise coordinate ϱ_* is defined as $\frac{d\varrho_*}{d\varrho} = \frac{A}{B}$, which yields

$$\varrho_* = \frac{1}{2} \sqrt{(Q_1 - Q_2)^2 + 4\varrho^2} + \frac{2Q_1 Q_2}{Q_1 + Q_2 - \sqrt{(Q_1 - Q_2)^2 + 4\varrho^2}} + (Q_1 + Q_2) \ln \left(Q_1 + Q_2 - \sqrt{(Q_1 - Q_2)^2 + 4\varrho^2} \right). \quad (\text{B5})$$

Finally, the evolution equation for $\sigma_{lm}(t, R)$ in tortoise coordinates is governed by

$$-\frac{\partial^2}{\partial t^2} \sigma_{lm} + \frac{\partial^2}{\partial \varrho_*^2} \sigma_{lm} - V_{\text{eff}}(\varrho) \sigma_{lm} = 0 \quad (\text{B6})$$

where the effective potential reads

$$V_{\text{eff}}(\varrho_*) = B^2 \left(\frac{l(l+1)}{\varrho^2} + \frac{1}{\varrho A^2} \left(\frac{\partial_\varrho B}{B} - \frac{\partial_\varrho A}{A} \right) \right). \quad (\text{B7})$$

We integrate Eq. (B6) numerically with an ingoing boundary condition at $\varrho_* \rightarrow -\infty$ (i.e., $\partial_t \sigma_{lm} \sim \partial_{\varrho_*} \sigma_{lm}$), and outgoing boundary condition at $\varrho_* \rightarrow \infty$ (i.e., $\partial_t \sigma_{lm} \sim -\partial_{\varrho_*} \sigma_{lm}$).

For the initial data for simulation we choose an instantaneously static Gaussian initial data,

$$\sigma_{lm}(0, \varrho_*) = a e^{-\left(\frac{\varrho_* - \varrho_0}{\sigma}\right)^2}, \quad (\text{B8})$$

$$\dot{\sigma}_{lm}(0, \varrho_*) = 0, \quad (\text{B9})$$

where a , ϱ_0 , and σ are the initial amplitude, position, and width of the Gaussian pulse, respectively. Note that our initial data are given for each (l, m) mode separately, since the modes are decoupled in the spherically symmetric background. We used a C++ numerical code to integrate Eq. (B6) using a 4th-order Runge-Kutta method, where spatial derivatives are evaluated by 4th-order finite difference.

2. Frequency-domain analysis

In the spherically-symmetric case it is straightforward to integrate the Klein-Gordon equation as an eigenvalue problem in order to compute directly the QNMs of the 4-charge BH.

In the particular case $Q_1 = Q_3$ and $Q_2 = Q_4$, Eq. (B6) reduces to a single ordinary differential equation in a Schroedinger-like form by assuming $\sigma_{lm} = \Psi(\varrho) e^{-i\omega t}$. Boundary conditions are then imposed ($\Psi \rightarrow e^{\pm i\omega \varrho_*}$ as $\varrho_* \rightarrow \pm\infty$) in order to compute the complex eigenvalues ω . In the general case, $r_H^4 = Q_1 Q_2 Q_3 Q_4$, the same method can be implemented by integrating the Klein-Gordon equation on the metric (7) using directly isotropic coordinates. By decomposing the field as $\Phi = \sum_{lm} \frac{F(r)}{r} Y_{lm} e^{-i\omega t}$, the Klein-Gordon equation takes a particularly simple form,

$$F'' + \left(\omega^2 \frac{(Q_1 + r)(Q_2 + r)(Q_3 + r)(Q_4 + r)}{r^4} - \frac{l(l+1)}{r^2} \right) F = 0. \quad (\text{B10})$$

In both cases, we computed the fundamental eigenfrequencies ω using an extended version of the direct integration method [107, 108]. Following [109], the QNMs can also be found ‘analytically’ by means of the ‘quantum’ Seiberg-Witten curve of $\mathcal{N} = 2$ super Yang-Mills theory with $SU(2)$ gauge group and RG-invariant scale Λ , coupled to $N_f = 2$ hypermultiplet doublets with masses $m_{1,2}$, after the identifications $\Lambda = +2i\hbar\omega r_H$, $m_1 = \frac{\hbar}{2} - 2i\hbar\omega \sum_A Q_A$, $m_2 = -\frac{\hbar}{2} - 2i\hbar\omega \sum_A \frac{r_H^2}{Q_A}$ and $u = \hbar^2 \left[(l + \frac{1}{2})^2 + \omega^2 (2r_H^2 - \sum_{A < B} Q_A Q_B) \right]$ [110].

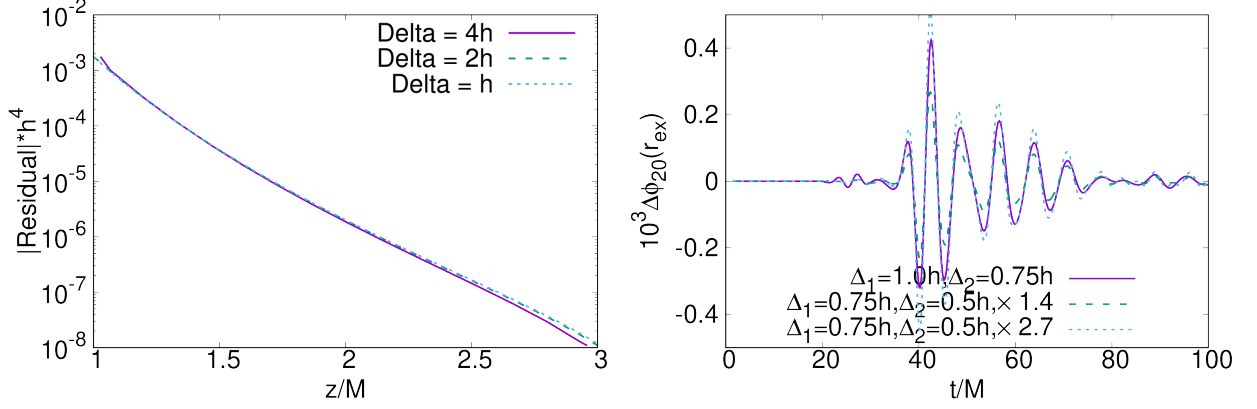


FIG. 8. Convergence of the residual of the differential form in Eq. (2) on z -axis for different resolutions h (left panel), and convergence analysis of the $l = 2, m = 0$ multipole mode of the scalar field at $r_{\text{ex}} = 20M/3$ (right panel).

Appendix C: Code tests and regularization procedure for fuzzball simulations

In this appendix we provide some technical details of our code. First of all, since microstate geometries are cumbersome, it is useful to check the analytical solutions. For this purpose, in the left panel of Fig. 8, we check the numerical convergence of residual in the differential form of Eq. (2) for different resolutions. The residual shows convergence with fourth-order accuracy, as expected since the spatial derivatives are evaluated by fourth-order-accurate finite difference stencils.

The right panel of Fig. 8 shows the evolution of the difference $\Delta \Phi_{20}$ between the field Φ_{20} at different resolutions Δ_1 and Δ_2 for a simulation starting from instantaneously static Gaussian initial data with $\sigma = 0.33M$ around the microstate scaling solution with $\kappa = 1$, and $L = 0.67M$. In our code, the spatial derivatives are approximated with fourth-order-accurate finite difference stencils, and the time integration is performed by a 4th-order Runge-Kutta method. The boundary of the mesh refinement boundary is interpolated with second order and fifth order accuracy in time and space. We define the convergence factor

$$Q_n = \frac{\Phi_{20}(\Delta_c) - \Phi_{20}(\Delta_m)}{\Phi_{20}(\Delta_m) - \Phi_{20}(\Delta_h)} = \frac{\Delta_c^n - \Delta_m^n}{\Delta_m^n - \Delta_h^n}, \quad (\text{C1})$$

where $\Delta_c = \Delta$, $\Delta_m = 0.75\Delta$, $\Delta_h = 0.5\Delta$, and n is expected convergence factor. As shown in the right panel of Fig. 8 the numerical evolutions converge between second- and third-order accuracy, as expected.

Finally, we discuss the regularization procedure to resolve the singularities of the 4-dimensional fuzzball geometry near the centers. In Fig. 9 we show an example of evolution for different values of the regularization parameter ϵ . While the initial response is independent of ϵ – since it is related to the properties of the photon sphere and not to the region near the centers – the behavior at later times depends on ϵ if the latter parameter is not sufficiently small. The two smallest values of ϵ give the same evolution for $l = 0$, whereas small finite- ϵ differences arise at late times for $l = 2$. Such differences become smaller as $\epsilon \rightarrow 0$ and, anyway, do not affect the overall structure of the signal. While using even smaller values of ϵ is computationally very expensive, our simulations indicate that for sufficiently small ϵ the evolution is regular and smoothly convergent.

Appendix D: Extra simulations

For completeness, in this appendix we consider two further examples of evolution in the axisymmetric microstate geometry with $\kappa_1 = \kappa_2 = \kappa_3 = \kappa$. In particular, in the left and right panels of Fig. 10 we consider the evolution of the $l = m = 2$ initial data on a solution with $\kappa = 2$ and $\kappa = 4$, respectively. The former case is interesting because, according to our geodesic analysis, it does not possess a light ring on the equatorial plane. Correspondingly, we do not observe echoes in this case since radiation confined near the equator cannot be efficiently trapped. On the other hand, the standard prompt ringdown shown in the left panel could be identified with a nontrivial geodesic structure outside the equatorial plane for this geometry. Finally, in the case $\kappa = 4$ there is an equatorial light ring which can efficiently confine radiation. Furthermore, in this case the fuzzball's throat is deeper and the associated echo delay

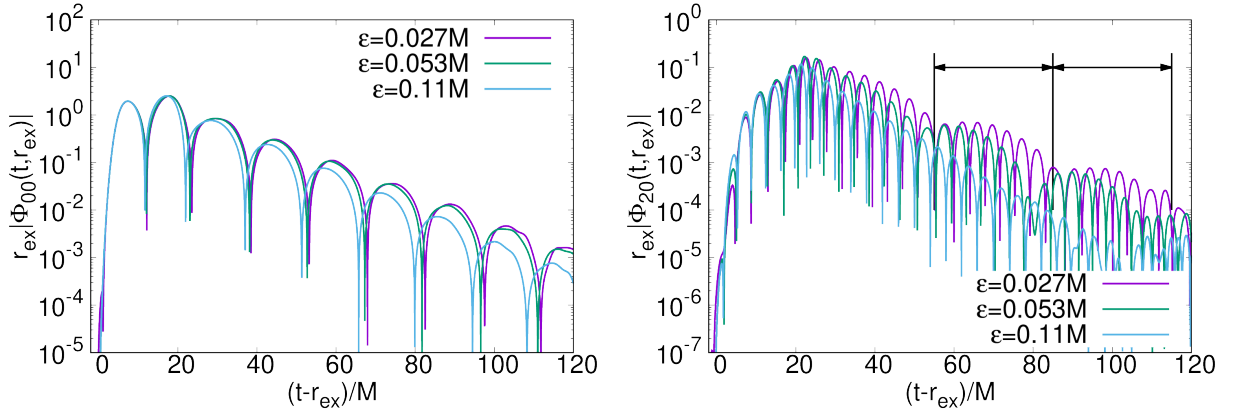


FIG. 9. Time evolution of $l = 0, m = 0$ (left panel) and $l = 2, m = 0$ (right panel) multipole modes of the scalar field around scaling solution with $L = 0.67M, \kappa = 1$. The evolution starts from instantaneously static spherically symmetric Gaussian profile Eq.23 with $\sigma = 0.33M$. The extraction radius is $r_{\text{ex}} = 20.0M$.

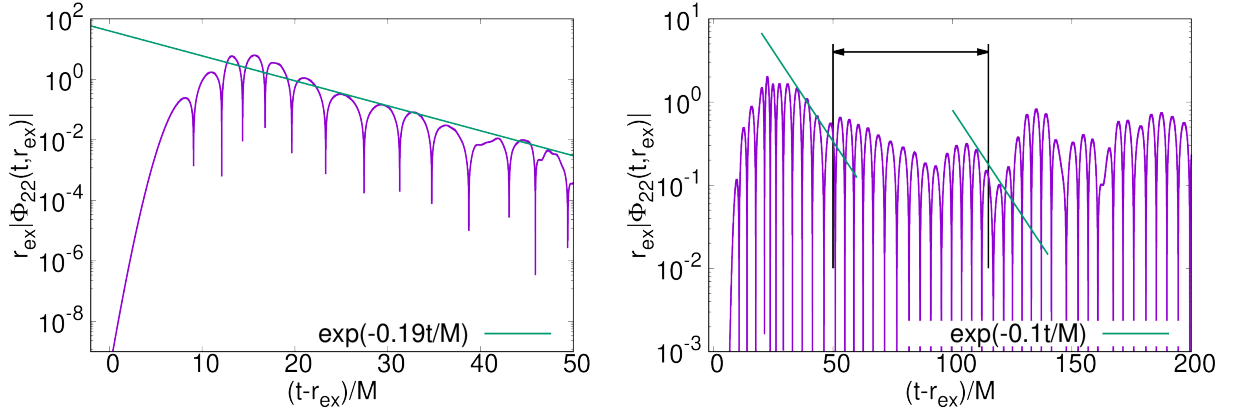


FIG. 10. Evolution of an instantaneously static $l = m = 2$ initial Gaussian profile on the axisymmetric microstate geometry with $\kappa_1 = \kappa_2 = \kappa_3 = \kappa$. Left panel: $\kappa = 2$; in this case the absence of an equatorial light ring implies the absence of echoes at late time. Right panel: $\kappa = 4$; also in this case the echo delay time is consistent with the geodesics prediction in Table I.

time is about $64M$ (see Table I). This is in qualitative agreement with the timescales shown in the right panel of Fig. 10, confirming that the physical picture drawn from the geodesics approximation is reliable.

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